

Percival Lowell

"the maximum speed [a molecule] may attain Clerk-Maxwell deduced from the doctrine of chances to be seven-fold the average. What may happen to one, must eventually happen to all."

Evidence of Ancient Escape

Earth

Xe isotopic fractionation, scarcity of **radiogenic** ^{129}Xe , ^{136}Xe

O_2 atmosphere (speculative, this)

Venus

dessication, high D/H ratio in rump water

Mars

Xe fractionation, extreme scarcity of **radiogenic** ^{129}Xe , ^{136}Xe

heavy D/H, $^{15}\text{N}/^{14}\text{N}$, $^{36}\text{Ar}/^{38}\text{Ar}$ ratios

heavy C, O

general absence of atmophiles

Io

Sulfur? follow the water?

Thermal Escape

“Jeans escape” = escape one hot atom or molecule at a time

escape is from an “exobase”

defined as scale height = mean free path

rule of thumb: scale height < Planet Radius/6

its relatively slow, but not necessarily unimportant

examples include modern H on Earth, Mars, Titan

it is extremely mass fractionating

it is not usually discussed for Io

“Jeans” Escape (what’s in a name)

1846 - Waterston suggested thermal escape from atmospheres. Paper rejected.

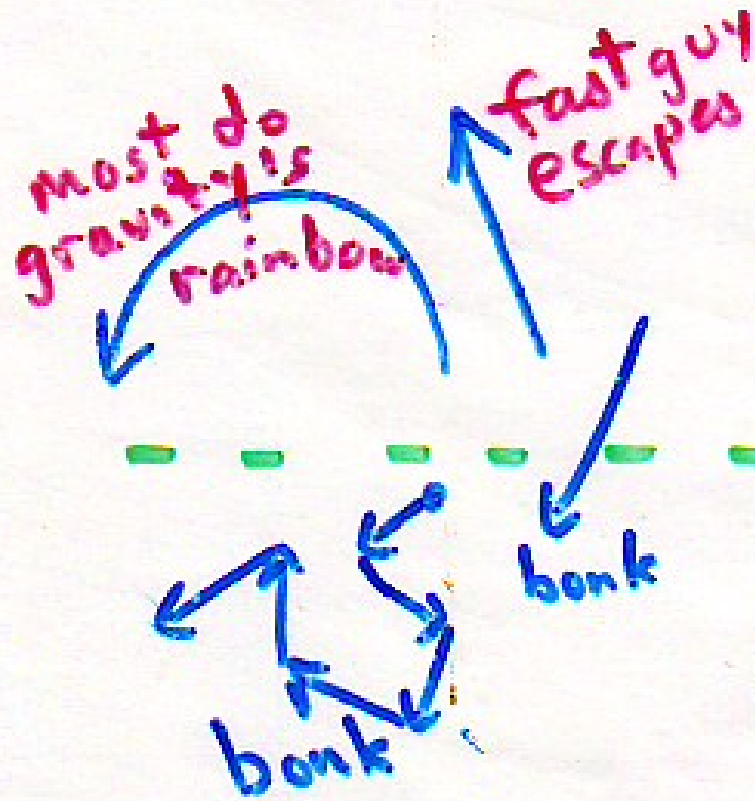
1868, 1898 - Stoney introduced concept of “**exosphere**,” rediscovered thermal escape.

1904 - Jeans reinvents exosphere and thermal escape. He ignores all previous work. Thermal escape becomes known as “**Jeans escape**.” IAU-approved.

Jeans escape describes the loss of atoms or molecules from an **exobase** that is treated like a surface

Exobase (aka "critical level")

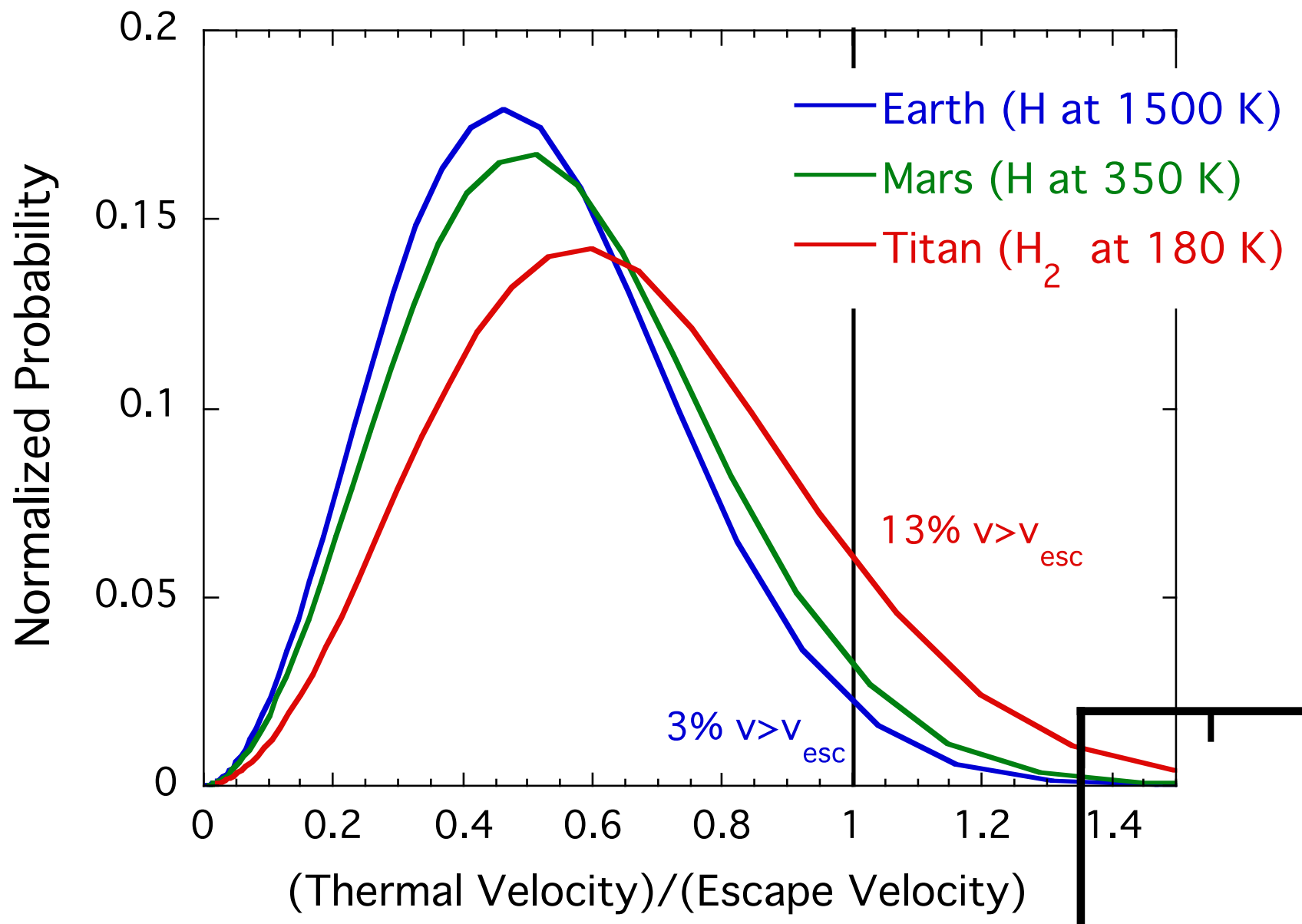
what goes up, may come down



collisionless
every molecule is just
another satellite

----- exobase

collisional



Thermal Escape continued

“hydrodynamic escape” = pressure-driven fluid flow

There is no exobase

its usually faster. Probably the only way to lose a Ganymede of H_2O

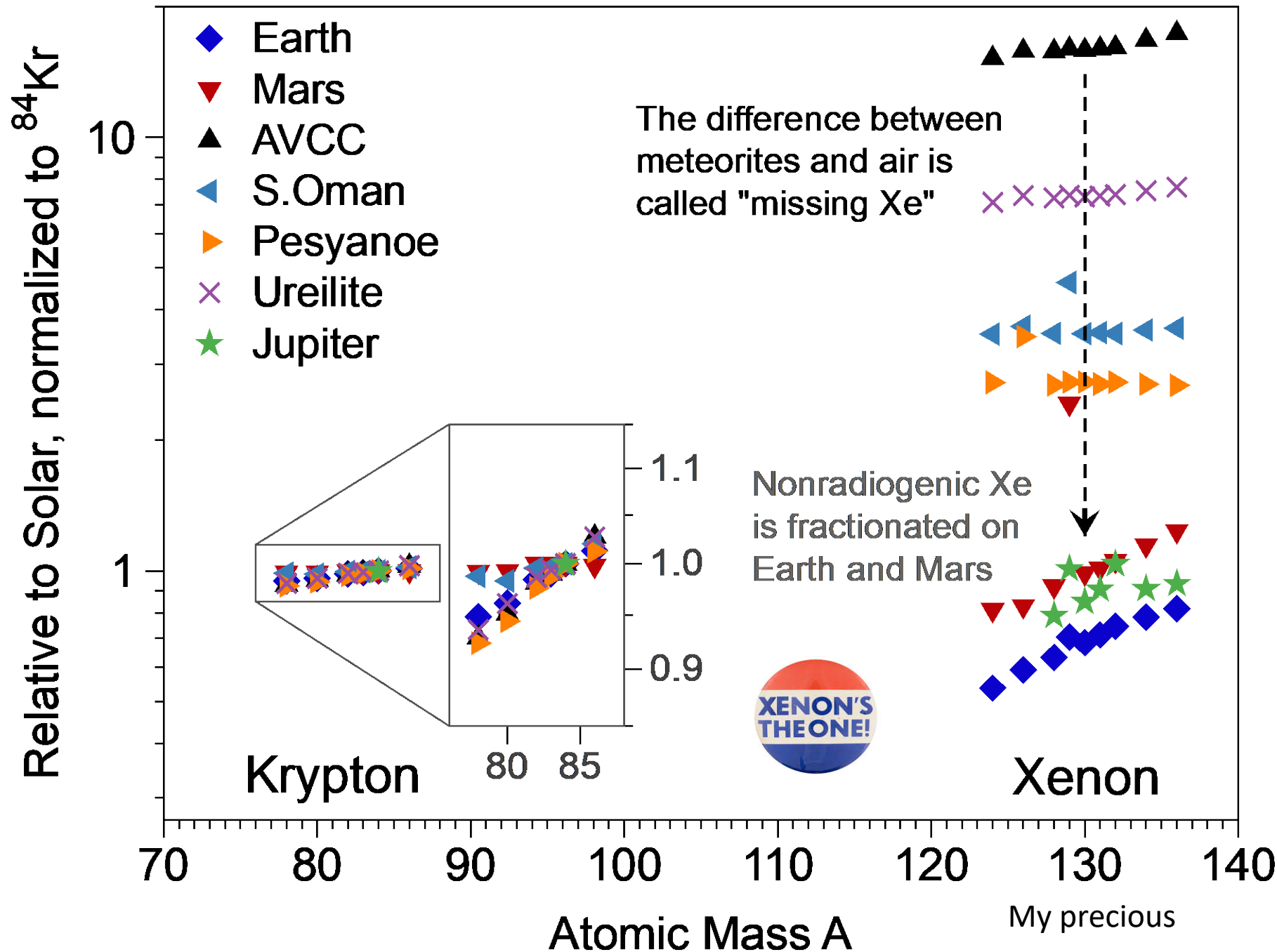
- solar wind; ancient Earth and Mars; possibly current Pluto
- doubtless ancient Venus

the canonical escape flux from Io is only 3-10 X smaller than would require a hydrodynamic description

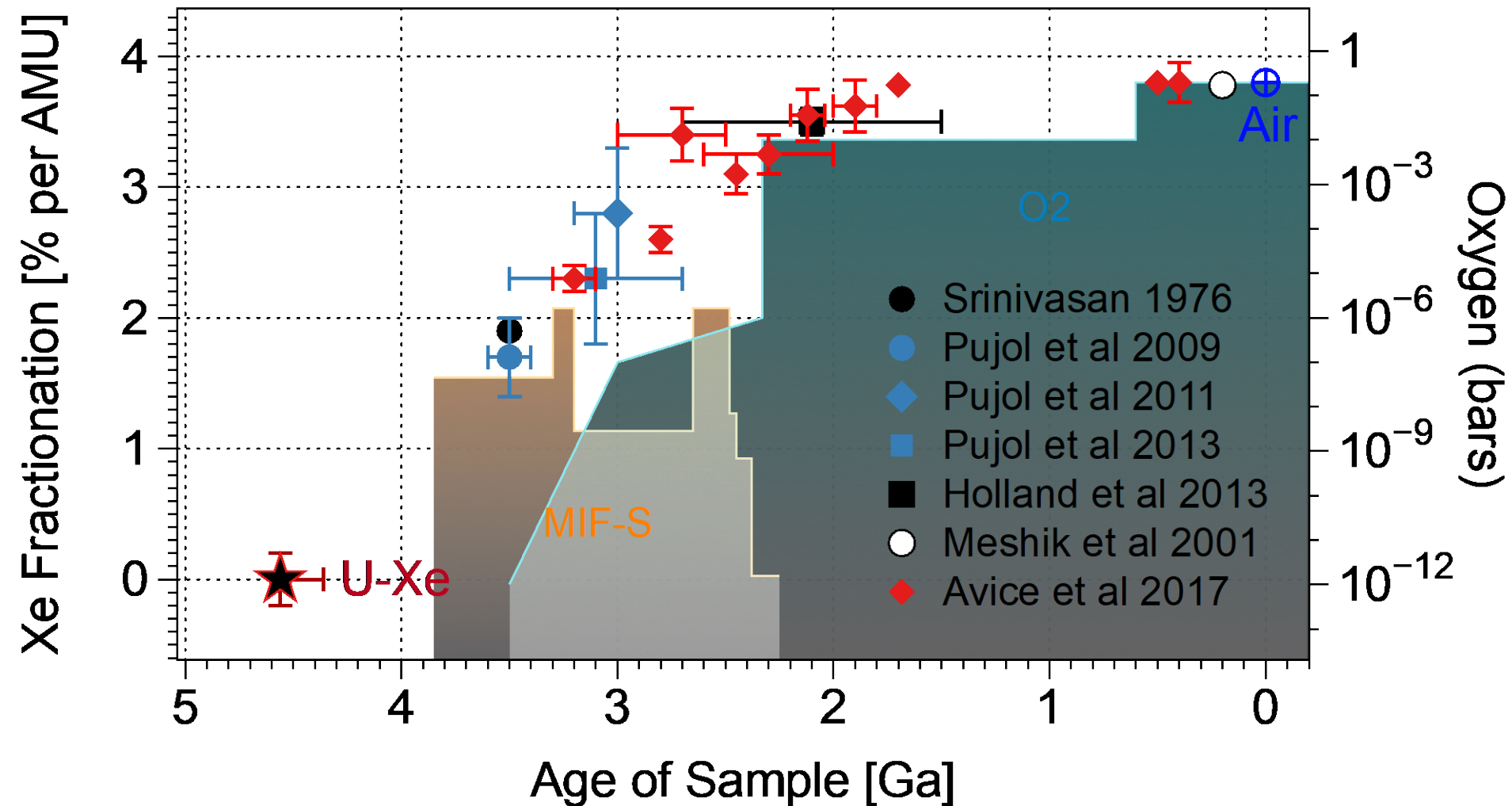
Fractionation in hydrodynamic escape occurs when the carrier gas drags heavier species into space. My favorite e.g. is Xenon.

Fractionation depends on the flux of the carrier gas

- if the flux much exceeds what is required to drag off the heavy gas, fractionation is modest
- if the flux is just barely sufficient to carry off the heavy gas, fractionations can be large with very little actual escape



Adapted from Guillaume Avice et al 2017



Nonthermal Escape (many many mechanisms)

most but not all are one atom at a time, and therefore most are slow

charge exchange ($\text{fast } X^+ + \text{slow } X > \text{fast } X + \text{slow } X^+$)

H and O on Earth, Venus, Mars.

Na, O, S, Cl on Io

dissociative recombination ($XY^+ + e^- > \text{fast } X + \text{fast } Y$)

N, O, C on Mars [Titan?]; **NaCl, SO, S₂, O₂ at Io**

solar wind sweeping, sputtering, ion pickup etc

Ar, N on Mars

S, O, Na, Cl on Io

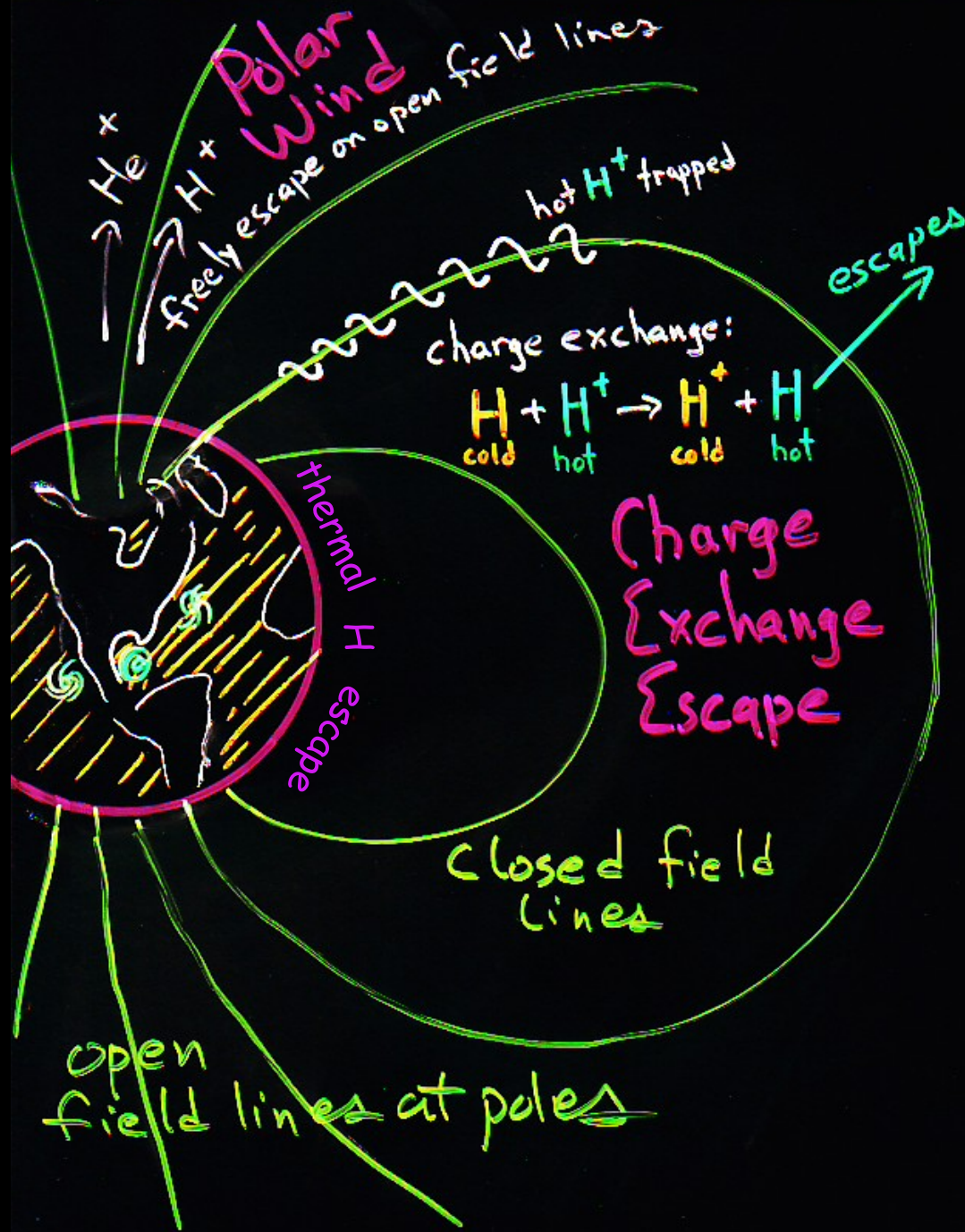
polar wind (electric field acceleration)

this is a kind of nonthermal hydrodynamic escape; possibly fast

He⁺ on Earth, H⁺ on Venus; H⁺, O⁺, N⁺, Xe⁺ on ancient Earth, Mars

probably something like this with ions and Io (the big current)

fractionating power varies



Diffusion limited escape – applies generally

limits how quickly one gas can diffuse through another*

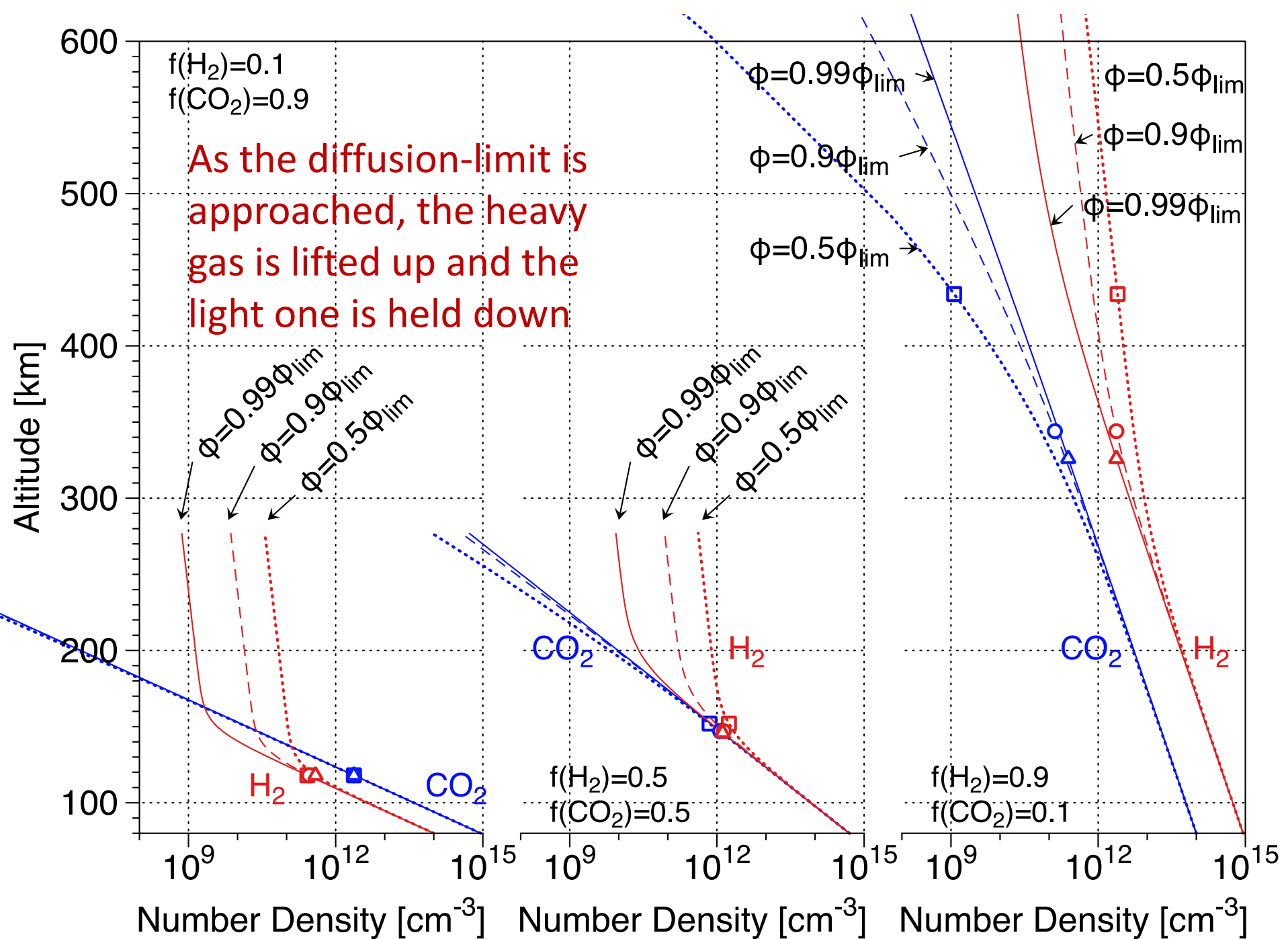
$$\phi_{\text{DL}} = b_{ij} \left(\frac{1}{H_i} - \frac{1}{H_j} \right) f(\text{H})$$

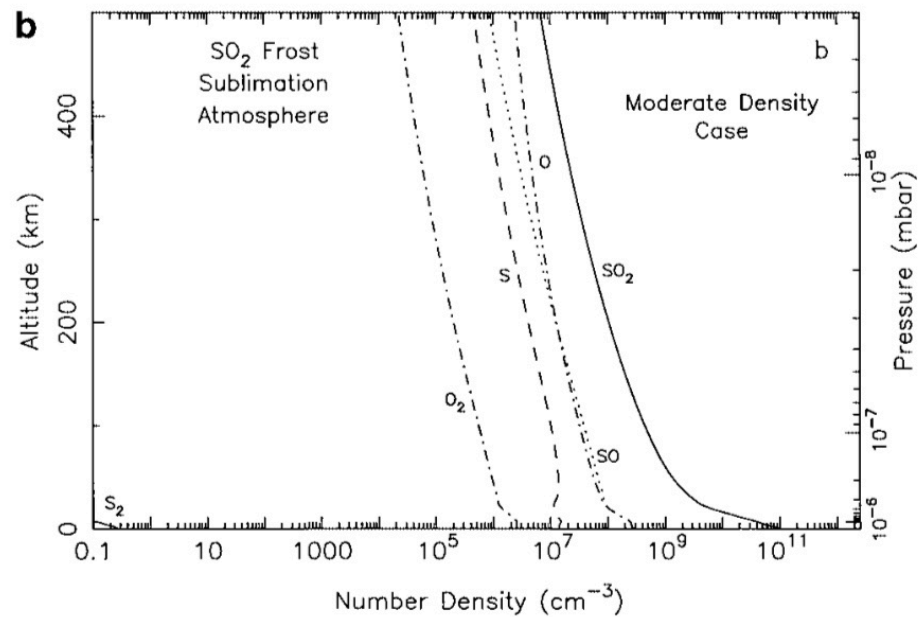
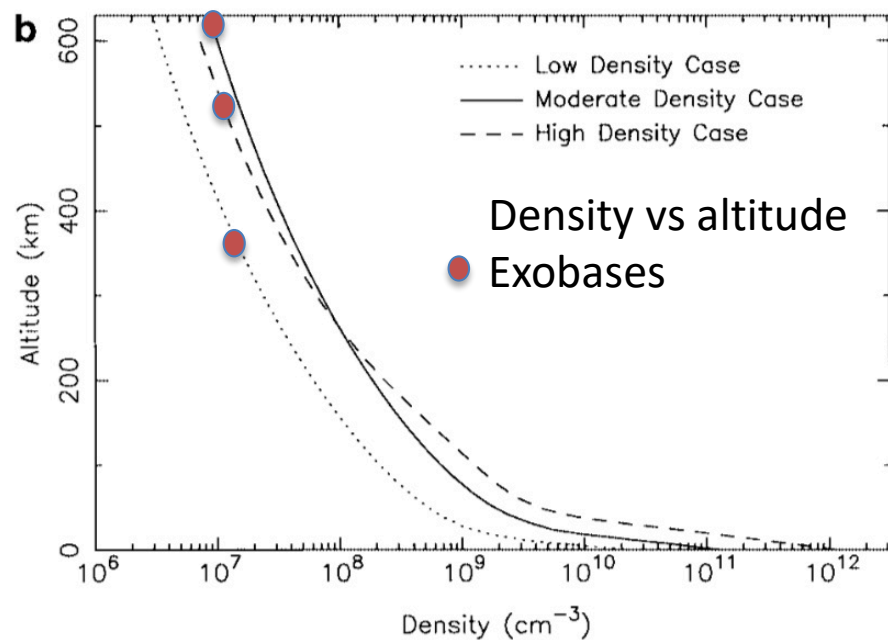
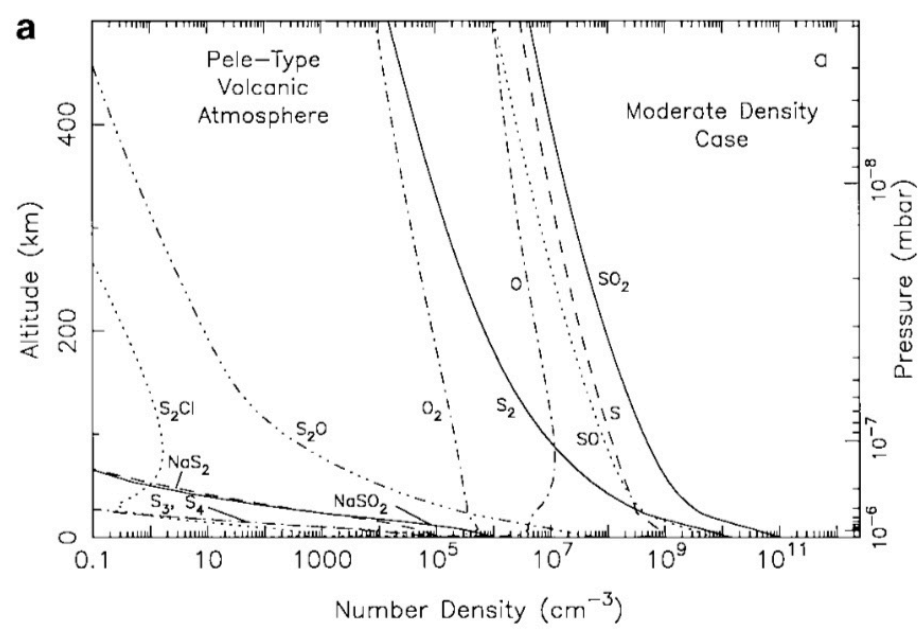
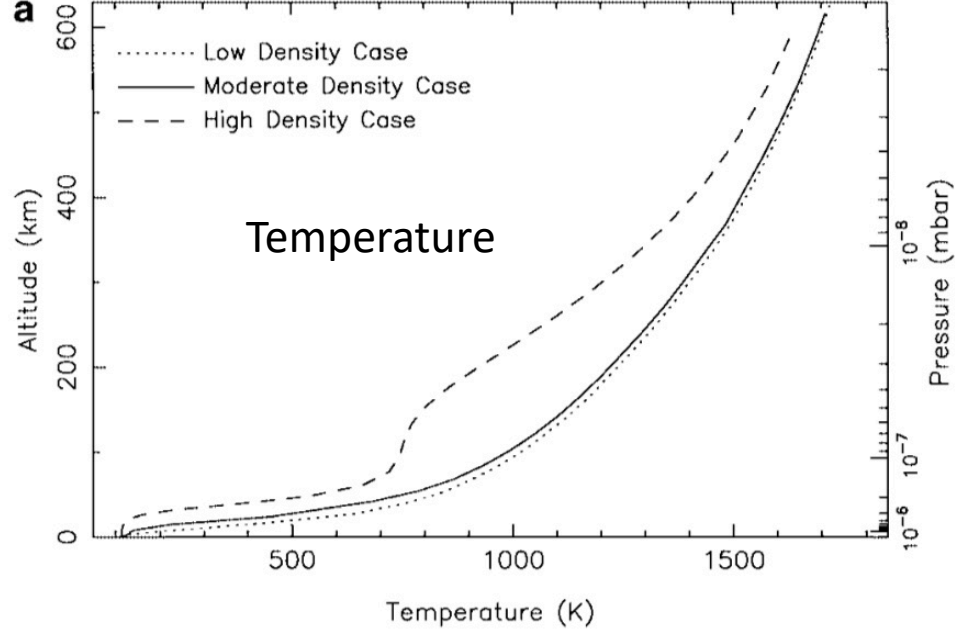
$f(\text{H})$ is the total hydrogen mixing ratio
and b_{ij} is the binary diffusion coefficient [$\text{cm}^{-1}\text{s}^{-1}$]

b_{ij} is the thermal velocity divided the collision cross section

Moses et al 2002 estimate diffusion limits for Na, Cl, S, O with respect to SO₂ for Io

*To do better requires phases separating – e.g., solids precipitating





reported escape fluxes ($\times 10^{26}$ per second)

canonical*

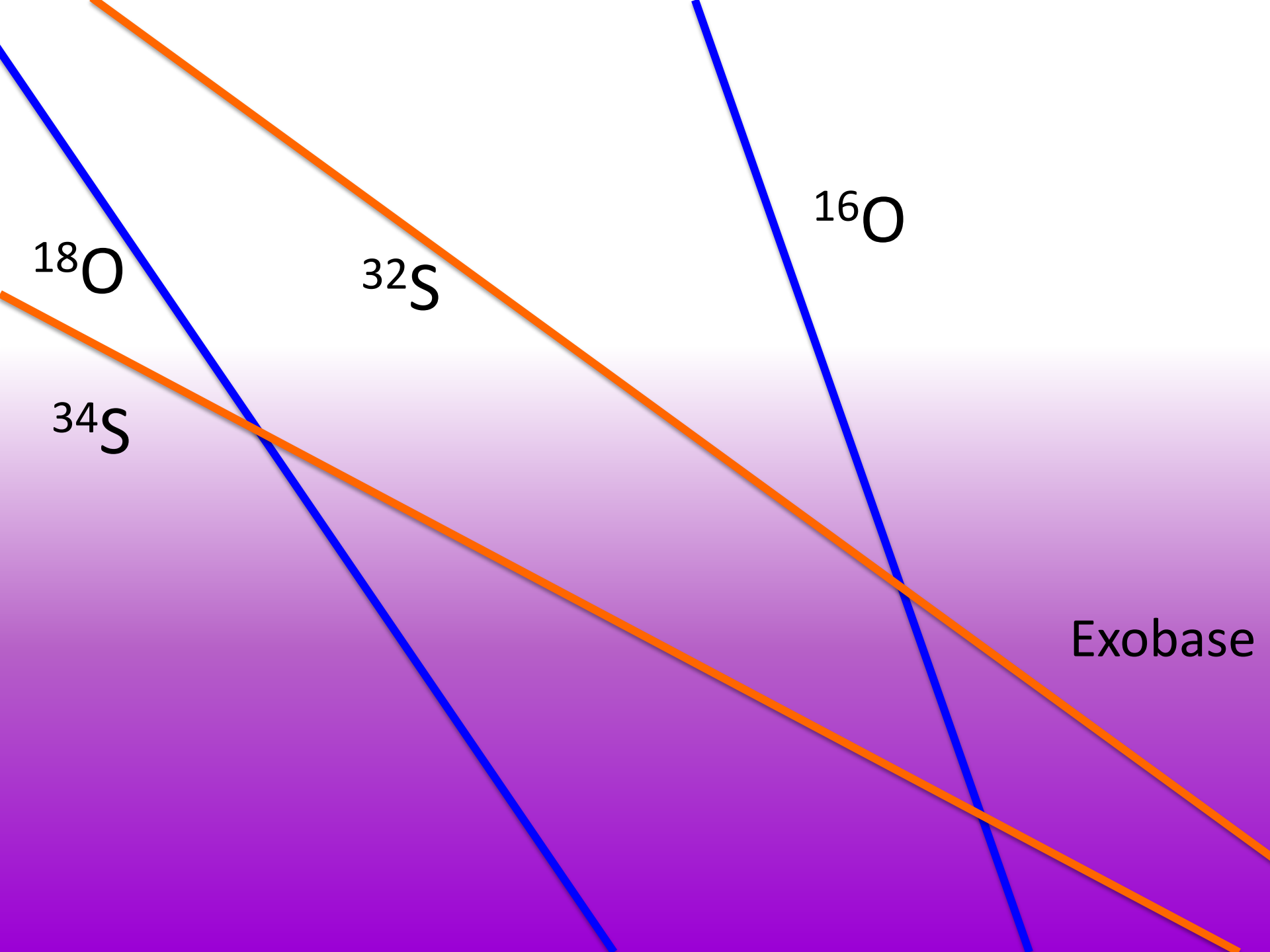
Moses et al 2002**

O	180	110
S	90	200
Na	3	50
NaCl	6	4
SO ₂	10	0
Cl		30

sum **1.1e6 g/s = 2 km of lost S over 4 Gyr**

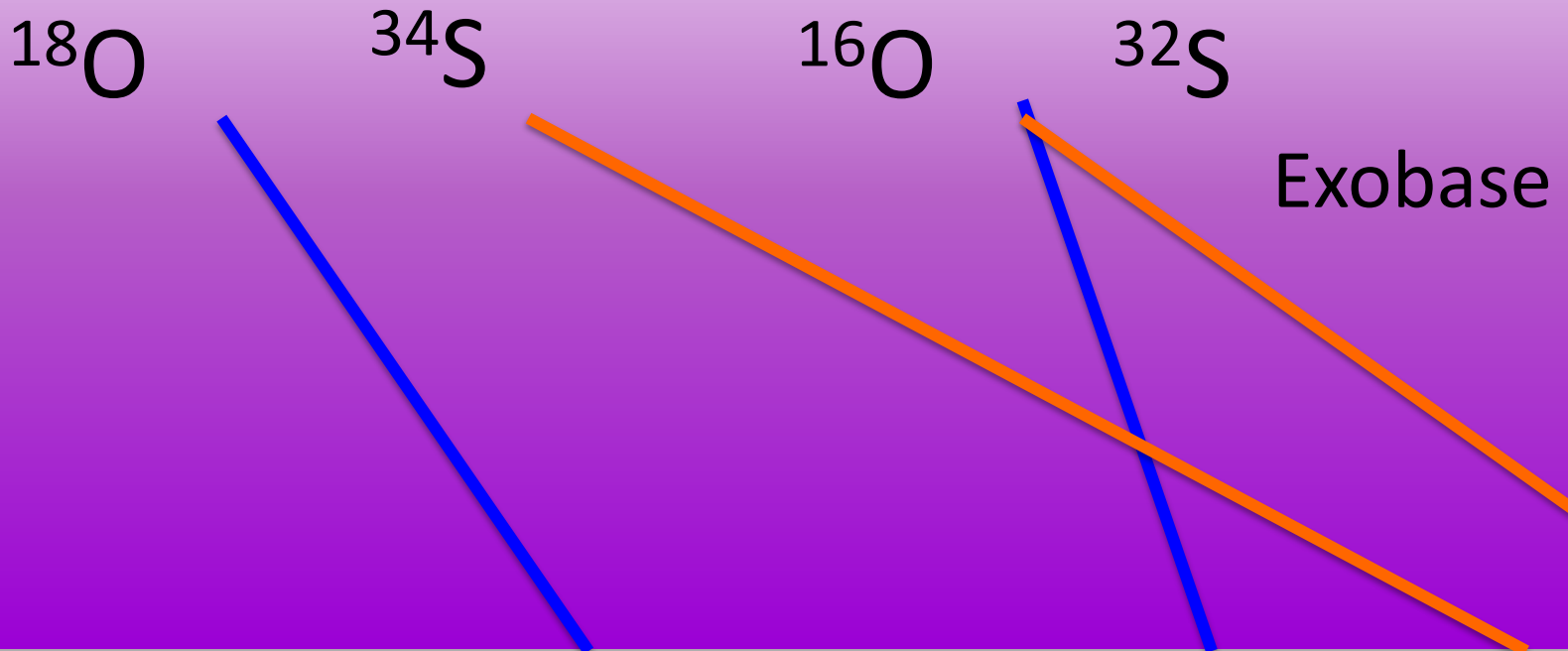
~~*scaled from Na.~~ mass flux based on mass-loading

diffusion-limiting flux, global **Pelean atm assumed



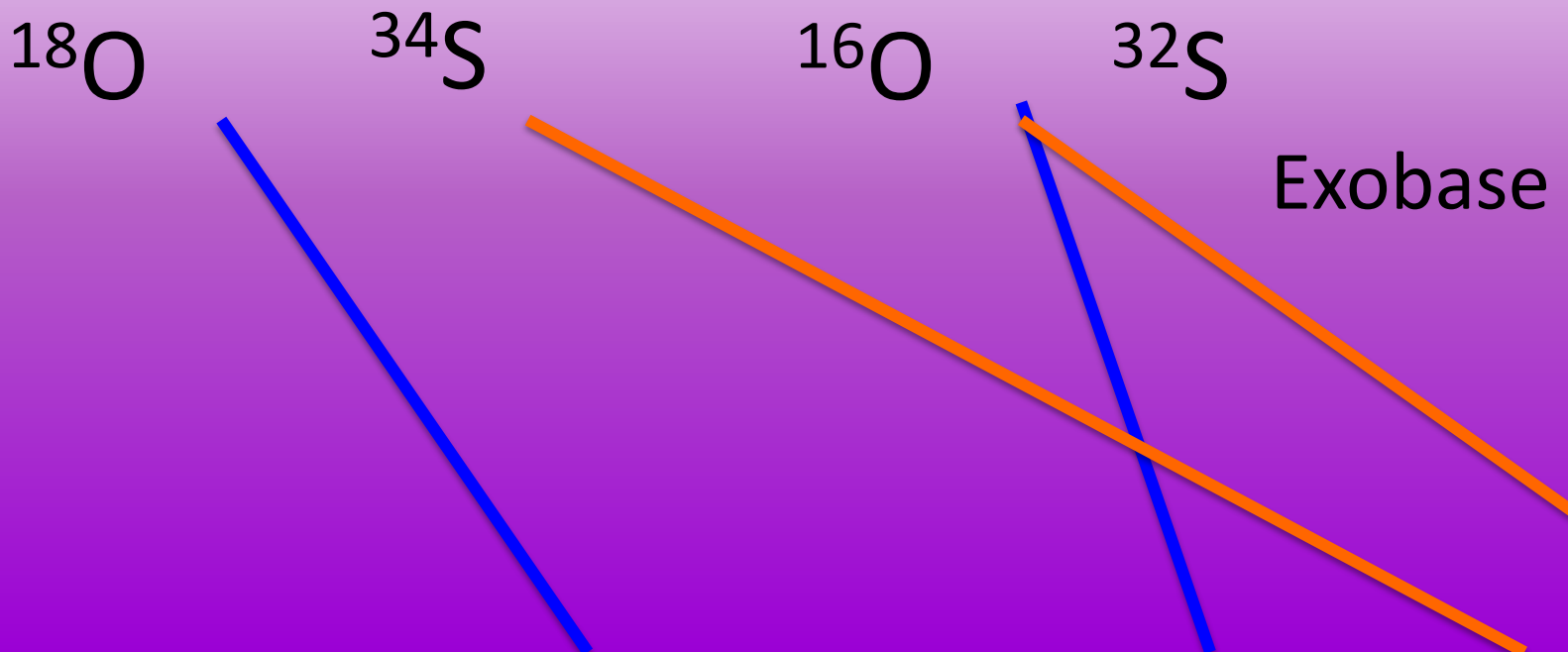
Sputtering picks off atoms ~~above~~ in general vicinity of the exobase

- which leaves the atmosphere enriched in heavier isotopes **and** enriched in **S** vs **O**



Ion pickup will prefer to pick off ions ~~above the exobase~~

- which goes the other way, favoring S^+ escape over O^+ escape because S is much easier to ionize. Na is much easier to ionize still.



some observations:

- escape takes place from the top of the atmosphere
no Ca, Fe, Si, Mg in the torus
- plumes enable escape of NaCl and products of S_n
escape of Na is probably restricted to plume tops
- two kinds of plumes:
Promethean, driven by SO_2 vapors
Pelean, driven by S_n vapors
- the atmosphere and plumes are substantially heated
solar UV alone would heat a plume by 1-4 K/s
particle heating is said to be bigger,
Joule heating is said to be bigger still

more speculations, mostly re isotope fractionation

- 2 km sulfur is 0.16% of Io's mass.
this seems like a lot to me – its 10X the sulfur of BSE
- it seems reasonable to expect that S and O escape stoichiometrically as SO_2 .

This is similar to what is inferred of H and O (H_2O) escaping from Mars and Venus.

We should expect feedbacks between escape, atmospheric chemistry, and volcanism to maintain this

- isotopic fractionation of S and O should be correlated
- isotopic fractionation is sensitive to the relationship between the exobase and the homopause.

My guess is that neither will be well-defined, in which case fractionation will not be extremely sensitive to mass